

The High-Density Ionized Gas in the Central Parsec of the Galaxy

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ABSTRACT

We report a study of the H30 α line emission at 1.3 mm from the region around Sgr A* made with the Submillimeter Array at a resolution of 2'' over a field of 60'' (2 parsec) and a velocity range of -360 to $+345$ km s⁻¹. This field encompasses most of the Galactic center's "minispiral." With an isothermal homogeneous HII model, we determined the physical conditions of the ionized gas at specific locations in the Northern and Eastern Arms from the H30 α line data along with Very Large Array data from the H92 α line at 3.6 cm and from the radio continuum emission at 1.3 cm. The typical electron density and kinetic temperature in the minispiral arms are $3\text{--}21 \times 10^4$ cm⁻³ and 5,000–13,000 K, respectively. The H30 α and H92 α line profiles are broadened due to the large velocity shear within and along the beam produced by dynamical motions in the strong gravitational field near Sgr A*. We constructed a 3D model of the minispiral using the orbital parameters derived under the assumptions that the gas flows are in Keplerian motion. The gas in the Eastern Arm appears to collide with the Northern Arm flow in the "Bar" region, which is located 0.1–0.2 parsec south of and behind Sgr A*. Finally, a total Lyman continuum flux of 3×10^{50} photons s⁻¹ is inferred from the assumption that the gas is photoionized and the ionizing photons for the high-density gas in the minispiral arms are from external sources, which is equivalent to ~ 250 O9-type zero-age-main-sequence stars.

Subject headings: Galaxy: center — ISM: individual (Sgr A) — ISM: kinematics and dynamics — ISM: recombination line — radio lines: ISM

1. INTRODUCTION

The Galactic center harbors a supermassive black hole (SMBH) with a mass of about $4.2 \times 10^6 M_\odot$ (Ghez et al. 2008; Gillessen et al. 2009) at the position of the radio source Sgr A*. The inner parsec of the Galactic center region contains

a rich cluster of stars, including at least 55 OB stars within the central 0.5 pc (Paumard et al. 2006). There are several theories to explain how these stars got there. In situ formation, as in giant molecular clouds, is improbable because of the stars' large velocity dispersion. The accretion sequence may or may not include the formation of an intermediate mass black hole in a super-dense stellar cluster that stabilizes the cluster against tidal disruption (Hansen & Milosavljevic 2003). The

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mystery of the formation of these stars is closely linked to the balance of mass inflow and outflow toward the Galactic center and its interaction with the source Sgr A*. Infall of interstellar material may lead to star formation up to a certain radius from the SMBH, and the subsequent feedback from stellar mass loss may suppress further star formation. Both phenomena may determine whether Sgr A* is in a quiescent state of the active galactic nuclei cycle (Loeb 2004).

Earlier radio recombination line (RRL) studies (Schwarz, Bregman, & van Gorkom 1989; Roberts & Goss 1993; Roberts, Yusef-Zadeh, & Goss 1996) revealed complex structures of ionized gas, including the “minispiral” and “Bar,” whose dynamics remain unclear. Parts of these structures have been modeled as gas orbiting Sgr A* (Schwarz, Bregman, & van Gorkom 1989; Sanders 1998; Paumard, Maillard, & Morris 2004; Vollmer & Duschl 2000; Liszt 2003; Muzic et al. 2007; Zhao et al. 2009) while other parts of the structures may be gravitationally unbound (Yusef-Zadeh, Roberts, & Biretta 1998). Another possibility is that some of the nearby stellar clusters (IRS 13 and IRS 16) have powerful winds that, with the help of gravitational focusing, create large-scale, unbound flows (e.g., Lutz, Krabbe, & Genzel 1993). There are a number of limitations in the existing spectral line observations of these structures. In the cm–radio images made with the Very Large Array (VLA),³ the recombination line strength is low, the line-to-continuum ratio is low, and the bandwidth is limited, making high-velocity line wings especially difficult to detect. In the infrared, the picture given by the emission lines from the ionized gas is subject to corrections for dust extinction of ~ 3 mag in the central parsec (Scoville et al. 2003; Buchholz, Schödel & Eckart 2009; Schödel et al. 2010). A few key questions about the ionized medium associated with Sgr A* are:

- 1) Is the ionized material related to stars or star formation?
- 2) Is the ionized material streaming freely in the central gravitational potential, or do other mechanisms such as stellar winds and shocks influence the dynamics?

³The Very Large Array (VLA) is operated by the National Radio Astronomy Observatory (NRAO). The NRAO is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

- 3) What are the column densities (or emission measures), densities, kinetic temperatures, and dynamics of the HII gas in the vicinity of Sgr A* itself?

In this paper, we present new observational results on the ionized gas in the central parsec of the Galaxy based on observations of the H30 α line made with the Submillimeter Array (SMA)⁴ and we interpret them in combination with previous H92 α and continuum measurements made with the VLA. Section 2 discusses the observations and data reduction. Section 3 presents our results and the modeling of the physical conditions and dynamics of the ionized gas in the vicinity of Sgr A*. Section 4 discusses the ionizing properties and possible links to the massive star formation in the central parsec, and Section 5 summarizes the results and conclusions. We assume a distance of 8 kpc to Sgr A* throughout the paper. At this distance, 1 pc corresponds to 24''.

2. OBSERVATIONS AND DATA REDUCTION

2.1. SMA H30 α Line Data

The H30 α line observations reported here were obtained with the SMA as part of a key project on the Galactic center in the period of 2006–2008. Our data comprise eleven tracks (Table 1). In all cases, the pointing center was the position of Sgr A*, and the half-power beam width at the 231.901 GHz frequency of the H30 α recombination line was 53'' (2 parsec at the distance of Galactic center). The 2 GHz bandwidth of the SMA correlator corresponds to a velocity range of ± 1300 km s^{−1}. Six tracks were taken in the compact array configuration, where the synthesized beam was 5.1'' \times 3.2''. Five higher-resolution tracks were also acquired, four in the extended configuration (1.5'' \times 1.1'' resolution) and one in the very extended configuration (0.7'' \times 0.4'').

Calibrations of the data were carried out in Miriad (Sault, Teuben, & Wright 1995) with the specific implementation for SMA data reduction.⁵

⁴The Submillimeter Array is a joint project between the Smithsonian Astrophysical Observatory and the Academia Sinica Institute of Astronomy and Astrophysics and is funded by the Smithsonian Institution and the Academia Sinica.

⁵<http://www.cfa.harvard.edu/sma/miriad>

The system temperature correction was made in post-processing to compensate for amplitude attenuation owing to the Earth’s atmosphere. Using the calibrators listed in Table 1, we determined the antenna-based bandpass solutions for each observing track, and the bandpass corrections were applied to the visibility data. The flux density scale was determined from observations of planets. Phase corrections were made using the point source model of Sgr A*, conveniently located at the center of the field, based on self-calibrations. The continuum emission was subtracted from the spectral data using the Miriad task UVLIN for both the point source (~ 3 Jy) and the extended emission (~ 3 Jy).

We constructed the H30 α line image cube with robust weighting ($R = 2$) (Briggs 1995). The synthesized FWHM beam was $1.9'' \times 1.5''$ (PA = 26°). In order to compare the SMA images with the VLA images of the H92 α line and 22 GHz continuum emission, the SMA image cube was convolved to a circular beam of $2''$. The intrinsic resolution of the correlator was 3.2 MHz, or 4.2 km s^{-1} . We smoothed the spectra to a resolution of 15 km s^{-1} with a Miriad task. For this paper, we selected the velocity range of -360 to $+345 \text{ km s}^{-1}$ for analysis. We found no significant emission beyond this range. The rms noise level of each channel image was 7 mJy. The image of the integrated H30 α line is shown in Figure 1 along with spectra in the directions of known infrared sources. Figure 2 shows the distribution of radial velocities determined from fitting the peak velocities of the H30 α line profile in the central $40''$ of the field. Notice the large velocity gradients along the Northern and Eastern Arms. The Western Arc does not show well in the SMA H30 α data because most of the line emission is below the SMA’s 4σ sensitivity limit. It is detected, however, in single-dish H30 α data taken with the IRAM 30m telescope.

2.2. VLA H92 α Line Data

The detailed calibrations and imaging of the H92 α line data have been described by Zhao et al. (2009). They used the line image cube convolved to a circular beam ($\theta_{\text{FWHM}} = 2''$), the same size as that of the H30 α line cube. Figure 3 shows a comparison of the spectra of the H92 α and H30 α lines toward selected regions containing IR sources. Table 2 lists the line fluxes and their ratios in these

spectra.

2.3. Continuum Data at 22 GHz

The continuum image at 22 GHz was obtained with the VLA high-resolution (A and B arrays) data set of 2005 described in Zhao et al. (2009) and the VLA archival data sets observed in C and D arrays during 2004 and 2005. The data reduction procedure described in Zhao et al. (2009) was followed. Sgr A* and the X-ray transient source (J174540.0290031) $3''$ south of Sgr A* were removed from the visibility data prior to imaging. The image of continuum emission was constructed from the visibility data with a long-baseline cutoff ($\leq 100 \text{ k}\lambda$) to achieve a synthesized FWHM beam of $1.95'' \times 1.41''$. The dirty image was cleaned with the Clark–Steer hybrid algorithm and convolved to a FWHM beam of $2''$, for comparison with other data at the same resolution.

2.4. IRAM H30 α Line Data

We also used additional H30 α data taken with the HERA multibeam receiver on the IRAM 30m telescope, with $11''$ beams, in order to calculate the global parameters for the region. Based on the intensity of integrated H30 α line emission in a lower-angular-resolution ($\sim 4''$) image constructed by combining the SMA data with the IRAM 30m data, we separated Sgr A West into two components: A, the bright minispiral features of the Northern and Eastern Arms ($\int S_{\text{H30}\alpha} dV \geq 5 \text{ Jy beam}^{-1} \text{ km s}^{-1}$) that trace the high-density ionized gas, and B, the residuals of Sgr A West after subtracting the emission from A. These results are in Table 2.

3. RESULTS

3.1. Physical Conditions of the Ionized Gas in the Central Parsec

The overall radio properties of the observed region (Sgr A West; $80'' \times 45''$) determined from the H30 α , H92 α , and 22-GHz continuum images are summarized in Table 2.

3.1.1. Distribution of H30 α Line Emission and IR Sources

Figures 1 and 2 show the H30 α line data in the central parsec. Figure 3 shows a com-

parison between the SMA H30 α and the VLA H92 α line spectra in the direction of selected IR sources. For optically thin gas in local thermodynamic equilibrium (LTE) with no pressure broadening, the ratio of line peak intensity is $R_{\text{H92}\alpha}^{\text{H30}\alpha} \equiv S_{\text{H30}\alpha}/S_{\text{H92}\alpha} \approx \nu_{\text{H30}\alpha}/\nu_{\text{H92}\alpha} \approx 28$. A detailed comparison of the H30 α line spectra from the SMA and the H92 α line spectra observed with the VLA yields line ratios in most of the regions that are close to the expected LTE ratio, suggesting that the ionized gas in the central parsec is optically thin and under LTE conditions. In some regions, however (e.g., IRS 1W, 10W, 16 and 33), the observed ratios show significant departure from the expected LTE ratio.

3.1.2. Distribution of T_e^*

The equivalent electron temperature, T_e^* , derived from the ratio of line-to-continuum emission under the assumption that the optically thin gas is in LTE, can be estimated from our measurements. In order to avoid contamination from both the dust emission, which becomes dominant at higher frequencies, and significant synchrotron emission at lower frequencies, we used the continuum data at 22 GHz, where the contributions from both the synchrotron and thermal dust emission are negligible relative to the free-free emission. Figures 4a and 4b show the integrated H30 α line image and radio continuum image at 22 GHz, respectively. We assume that the continuum emission at 22 GHz is entirely due to optically thin free-free emission and scale its intensity to 231.9 GHz, according to the power law $\nu^{-0.1}$, i.e., a factor of 0.79. Using Equation (1) below (Wilson, Rohlfs, & Hüttemeister 2009), we determined the value of T_e^* ,

$$\frac{T_e^*}{K} = \left[\frac{6985}{a(\nu, T_e)} \left(\frac{\nu}{\text{GHz}} \right)^{1.1} \frac{1}{1 + N(\text{He})/N(\text{H})} \left(\frac{S_C}{S_L \Delta V_{\text{FWHM}}} \right) \right]^{0.87}, \quad (1)$$

where ν is the frequency; S_L is the line flux density; S_C is the continuum flux density; ΔV_{FWHM} is the line width (FWHM), and only Doppler broadening is involved; $N(\text{He})/N(\text{H})$ is the number density ratio, which we take to be 0.08; and

$\alpha(\nu, T_e) \sim 0.97$ (at $\nu = 22$ GHz and T_e , the electron temperature, $\sim 10^4$ K) is the correction to the Mezger & Henderson (1967) power-law approximation. The angular distribution of T_e^* is shown in Figure 4c. T_e^* varies from 7,000 K to 10,000 K in most regions of Sgr A West. However, in the Bar region (e.g., IRS 2L, IRS 12N, IRS 13E, IRS 33), 3'' SE of Sgr A*, T_e^* increases to $\sim 15,000$ K. The significant increase of T_e^* in the Bar has also been derived from line observations of H92 α (Roberts & Goss 1993) and H66 α (Schwarz, Bregman, & van Gorkom 1989).

3.1.3. Isothermal Homogeneous HII Model

Since the observed quantities of line and continuum intensities depend on the density, temperature, and path length of HII gas, we can construct a model and fit the data to it over a wide frequency range in order to determine the parameters of the ionized gas. In particular, with the high-resolution data obtained from interferometer observations, uncertainties due to variations in structure and the physical parameters across the source are substantially reduced. Thus, the results from such an analysis become more reliable. Models for RRL emission from the nuclear region of external galaxies have been discussed by Puxley et al. (1991); Anantharamaiah et al. (1993, 2000); Zhao et al. (1996, 1997); Phookun, Anantharamaiah, & Goss (1998); and Rodriguez-Rico et al. (2004). They found that the main constraints for the models are the integrated RRL strength at multiple frequencies and the observed radio continuum spectrum. The observed quantities (line and continuum intensities) depend in a nonlinear way on the distribution of T_e and n_e , the electron density, both along the line of sight and across the telescope beam. In the cases of the extragalactic nuclear regions, at a given frequency, the line emission often arises from HII components under the physical conditions that are particularly favorable at this particular frequency. Usually, multiple collections of HII regions with different physical parameters are required to fit the observations.

Both the SMA and VLA data achieve angular resolution (2'', ~ 0.1 pc) adequate to resolve the ~ 1 pc minispiral structure. The effects due to changes of physical parameters across the source are therefore less significant. However, the electron temperature and density could possi-

bly change along the line of sight; T_e and n_e derived from the simple isothermal and homogeneous model correspond, therefore, to average physical properties of the HII gas in the telescope beam. If the HII gas in the beam along the path length is isothermal and homogeneous, then the line and continuum flux densities from a region subtending a solid angle Ω are given by (e.g., Shaver (1975)),

$$S_L = \frac{2k\nu^2}{c^2} \Omega T_e \left[\left(\frac{\tau_L/\beta_n + \tau_C}{\tau_L + \tau_C} \right) \left(1 - e^{-(\tau_L + \tau_C)} \right) - (1 - e^{-\tau_C}) \right], \quad (2)$$

and

$$S_C = \frac{2k\nu^2}{c^2} \Omega T_e (1 - e^{-\tau_C}), \quad (3)$$

where k is Boltzmann's constant, c is the speed of light, and the line (τ_L) and continuum (τ_C) optical depths are given by

$$\tau_L \approx 575 b_n \beta_n \left(\frac{\nu}{\text{GHz}} \right)^{-1} \left(\frac{n_e}{\text{cm}^{-3}} \right)^2 \left(\frac{L f_V}{\text{pc}} \right) \left(\frac{T_e}{\text{K}} \right)^{-5/2} \left(\frac{\Delta V_D}{\text{km s}^{-1}} \right)^{-1} \left(1 + 1.48 \frac{\Delta V_P}{\Delta V_D} \right)^{-1}, \quad (4)$$

and

$$\tau_C \approx 0.08235 \left(\frac{n_e}{\text{cm}^{-3}} \right)^2 \left(\frac{L f_V}{\text{pc}} \right) \left(\frac{\nu}{\text{GHz}} \right)^{-2.1} \left(\frac{T_e}{\text{K}} \right)^{-1.35} a(\nu, T_e), \quad (5)$$

respectively. b_n and β_n are the population departure coefficients; n_e is the electron density; ΔV_D and ΔV_P are the FWHM Doppler and pressure line widths in km s^{-1} , respectively; L is the path length; and f_V is the volume filling factor. Note that the last factor in Equation (4) provides a convenient approximation for the transition between the pure Doppler and pure pressure-broadened cases.

The Doppler broadening has a thermal and turbulent component and is given by

$$\Delta V_D = \sqrt{\frac{8 \ln(2) k T_e}{M_H} + (\Delta V_t)^2}, \quad (6)$$

where M_H is the mass of hydrogen, and ΔV_t is the turbulent width, which here includes microscopic turbulence and dynamical velocity gradients. In the central parsec, the thermal width ($\Delta V_{th} = \sqrt{\frac{8 \ln(2) k T_e}{M_H}}$) is less than the turbulent width (ΔV_t). The line width due to pressure broadening, computed by Brocklehurst & Leeman (1971), is given by

$$\Delta V_P = 1.7 \times 10^{-18} \frac{n_e}{T_e^{0.1}} N^{7.4} \text{ km s}^{-1}, \quad (7)$$

where N is the principal quantum number of the transition. The total line width is approximately the quadratic sum of ΔV_D and ΔV_P ,

$$\Delta V_{FWHM} \cong \sqrt{(\Delta V_D)^2 + (\Delta V_P)^2}. \quad (8)$$

ΔV_{FWHM} is an observable quantity. We solved for the parameters T_e , n_e , ΔV_t , and $L f_V$ for each point in the image in the following way. Since $\Delta V_P \ll \Delta V_{th} < \Delta V_t$ for reasonable values of T_e and n_e , we assume initial values of T_e and n_e , and solve for ΔV_t via Equations (6)–(8). We computed the values for b_n and β_n for the assumed values of T_e and n_e using the NLTE code of Gordon & Sorooshenko (2002), which followed the analysis of Walmsley (1990). We then formed the flux density ratios $R_{22\text{GHz}}^{H30\alpha} \equiv S_{H30\alpha}/S_{22\text{GHz}}$, $R_{22\text{GHz}}^{H92\alpha} \equiv S_{H92\alpha}/S_{22\text{GHz}}$, and $R_{92\alpha}^{H30\alpha} \equiv S_{H30\alpha}/S_{H92\alpha}$. The use of these ratios eliminates the solid angle Ω in Equations (2) and (3). At each point in the image, we used these ratios to solve for T_e , n_e , and $L f_V$. We then iterated the process, i.e., solved for a new estimate of ΔV_t with these values of T_e and n_e . The process converged in a few iterations.

The parameters derived for the line of sight for each of the regions are summarized in Table 3. Instead of $L f_V$, we report the emission measure ($EM = n_e^2 L f_V$). The ratio T_e/T_e^* is given in column 6, which mainly corrects for the effect due to underpopulated electrons at the level $N = 30$ as compared to that expected under the LTE; the ratio $\Delta V_t/\Delta V_{FWHM}$ in column 7; and $\Delta V_P^{H92\alpha}/\Delta V_D^{TH}$, the ratio of the H92 α line width of pressure broadening to that of thermal Doppler broadening, in column 8. The pressure broadening term ($\Delta V_P^{H30\alpha}$) is negligible.

The uncertainties in the physical parameters mainly come from the uncertainties in the measurements. Although the relation between the

physical parameters and observable quantities is nonlinear, we assess the errors through linear approximation. We use the optically thin assumption, which appears to be valid at the frequencies used in this paper. The fractional uncertainty in T_e , which is proportional to the fractional uncertainties of the measured ratios, $R_{22\text{GHz}}^{\text{H}30\alpha}$ and $R_{\text{H}92\alpha}^{\text{H}30\alpha}$, lies in the range from 7% for IRS 1W to 39% for IRS 12N. The fractional uncertainty in n_e is dominated by those of observed line ratios $R_{\text{H}92\alpha}^{\text{H}30\alpha}$ and ranges from 6% for IRS 1W to 36% for IRS 12N.

The variations of the H30 α line intensities reflect the changes of excitation conditions in the region. Our model fitting suggests that, given similar solutions in n_e and T_e , the large difference in the H30 α line intensity between the two extreme regions (IRS 1W and IRS 7) is mainly due to the difference of the emission measure (EM) or the path length along the line of sight (see Table 3).

In short, the line ratio of H30 α to H92 α varies from 15 ± 4 to 32 ± 4 . The derived mean values of the line optical depth are $\bar{\tau}_L(\text{H}30\alpha) = -0.001$ and $\bar{\tau}_L(\text{H}92\alpha) = -0.006$ with standard deviations of $\sigma_{\tau_L(\text{H}30\alpha)} = 0.0004$ and $\sigma_{\tau_L(\text{H}92\alpha)} = 0.005$ from the corresponding emission regions. The free-free optical depth at the H30 α line frequency $\tau_C(231.9 \text{ GHz})$ varies from 1×10^{-4} to 1×10^{-5} while $\tau_C(8.3 \text{ GHz})$ is in the range between 0.16 and 0.01. Because of the small optical depths in both line and continuum, the non LTE effects in the minispiral, in general, are not critical.

There is a diffuse nonthermal synchrotron component in the central parsec region (Ekers et al. 1983). Based on the flux density of 22 Jy at 20 cm in the central $40'' \times 70''$ from Ekers et al. (1983), we estimated that the contribution to the continuum flux density in a $2''$ beam is ~ 10 mJy under the assumption that the nonthermal spectral index $\alpha \sim -0.5$ ($S_\nu \propto \nu^\alpha$) and the emission are evenly distributed. The optical depth of the H92 α line is small ($\tau_L \sim -0.01$). The contribution from the stimulated H92 α emission of ~ 0.1 mJy is insignificant in the regions discussed in this paper. The contribution from the stimulated line emission by the dust emission at 232 GHz is negligible because of the small optical depth of the H30 α line.

The observed line widths of the HII gas in the minispiral arms at the Galactic center are dominated by ΔV_t , which is consistent with velocity gradients due to gas motion around Sgr A* or the

SMBH being primarily responsible for the broadening in the line profiles observed in the $2''$ beam.

3.1.4. Northern and Eastern Arms

The ionized gas of the high-density components in the Northern and Eastern Arms has a mean kinetic temperature of $\bar{T}_e \approx 6 \times 10^3 \text{ K}$ and an electron density of n_e in the range of 10^{4-5} cm^{-3} , which is consistent with the values (10^4 cm^{-3}) derived from Pa α line emission at $1.87 \mu\text{m}$ (Scoville et al. 2003) for a given uncertainty in extinction corrections in the near IR. The HII gas properties in the Northern and Eastern Arms are similar to these in HII regions around young and/or evolved massive stars.

3.1.5. The Bar

We confirm that the ionized gas in the Bar (e.g., IRS 2L, IRS 12N, IRS 13E, and IRS 33 in Table 3) appears to have a higher kinetic temperature in the range of $8.5 \times 10^3 \text{ K} \leq T_e \leq 1.3 \times 10^4 \text{ K}$, as suggested by earlier RRL observations (Roberts & Goss 1993). The electron density shows a range from a few times 10^4 to a few times 10^5 cm^{-3} .

The heating process may be complicated in the Bar because of complex gas interactions. There are several energy sources in the region. First, strong winds from the central stellar clusters such as IRS 16 and IRS 13 could provide kinetic energy to compress the ionized gas in this region. The heating due to the interaction between the winds and the ionized gas in the Bar likely occurs in the regions localized to the cluster sources.

In addition, the well-known minicavity (marked with the gray circle in Figures 4a and 4b), a region depressed of radio continuum, is located $3.5''$ SW of Sgr A* (Yusef-Zadeh, Morris, & Ekers 1989). This region is likely created by a fast wind outflow from the sources in the vicinity sweeping up the gas in the orbiting ionized streams, as suggested by the morphology of a bright [FeIII] line in the infrared (Lutz, Krabbe, & Genzel 1993). However, the minicavity is a few arcsec displaced from the high-kinetic-temperature region, which in fact is best seen as a depression associated with H30 α line emission. The minicavity can also be explained by Melia, Coker, & Yusef-Zadeh (1996) as a downstream, focused flow due to the accretion by the

SMBH at Sgr A*.

Alternatively, the high temperature in the Bar could result from a collision between two ionized orbiting flows, as suggested in the dynamical model of Zhao et al. (2009). A computation using their Keplerian orbital model (see Section 3.2) shows that the specific kinetic energy reaches a maximum (see Figure 4d) in the Bar region where high kinetic temperatures (see Figure 4c) are observed. The distribution of specific kinetic energy (Figure 4d) shows a substantial increase from 2×10^{14} erg g⁻¹ at a region near IRS 5 in the Northern Arm to 12×10^{14} erg g⁻¹ at IRS 33 in the Bar. However, we note that a high kinetic energy of the orbiting flow does not have to correspond to a high kinetic temperature; the ionized flows can release their kinetic energy to shocks if the ionized flows collide with each other. If they collide with the strong stellar winds from the clusters of stars in the Bar area, then the gas could be heated by the shocks to higher temperatures. In addition, the model also shows that the ionized flows of the Northern and Eastern Arms do collide at the region with high kinetic temperatures (see Figure 21d of Zhao et al. (2009) and the 3D structure of the minispiral arms in Figure 5). The observed high kinetic temperatures in the Bar suggest that the ionized gas could be further heated by the shocks generated from the collision.

3.2. Kinematics and Dynamics

Kinematics of the ionized gas in both the Northern Arm and Western Arc have been explained as Keplerian circular motion of a circumnuclear ring (Schwarz, Bregman, & van Gorkom 1989; Lacy, Achtermann, & Serabyn 1991; Vollmer & Duschl 2000; Liszt 2003). Based on the VLA observations of H92 α , Roberts & Goss (1993) found that the ionized ring consists of three separated kinematic components, the Western Arc, Northern Arm, and Bar. The radial velocities of the Western Arc can be fitted with a Keplerian model with circular motion. Using the observations of the Brackett- γ line, Paumard, Maillard, & Morris (2004) found that the kinematics of the ionized gas in the Northern Arm can be explained as a bundle of elliptical orbits. In addition, the Eastern Arm may be interrupted by a high-velocity component “tip”. Also, the Bar component extends to the IRS6/34 region which is affected by two

other kinematic features, the “Western Bridge” and the “Bar Overlay.” The VLA observations of the H92 α line suggest that the loci of the three minispiral arms can be fitted with three families of elliptical orbits; the orbits of both the Northern and Eastern Arms have high eccentricities while the Western Arc is in nearly circular motion (Zhao et al. 2009). The VLA measurements of the proper motions and the radial velocity resolved the degeneracies in the inclination angles, suggesting that the two ionized flows from the Northern and Eastern Arms collide in the Bar region.

On the other hand, Muzic et al. (2007) suggested that the non-Keplerian motions might be important in the kinematics of the filaments observed near Sgr A*. In addition, Vollmer & Duschl (2000) noticed that a substantial amount of velocity dispersion along with additional inwards radial motion is required in order to fit the observed velocities of the ionized flows with their Keplerian circular-orbit model.

3.2.1. 3D View of the Minispiral Structure and the Circumnuclear Disk

Using the Keplerian model described in Zhao et al. (2009) along with the orbital parameters derived from fitting the VLA H92 α data and proper motion measurements, we recomputed the 3D locations of the three ionized arms. Figure 5 shows a 3D visualization of the three ionized flows in the Northern, Eastern, and Western Arms (Arc), which illustrates that the Northern Arm and Western Arc are nearly coplanar based on the similar longitudes of their ascending nodes (Ω) and inclination angles (i) derived from fitting the loci of the ionized flows. The mean values of the two orbital angles, $\bar{\Omega} = 71^\circ \pm 6^\circ$ and $\bar{i} = 119^\circ \pm 3^\circ$, are inferred for the two coplanar orbits (Northern Arm and Western Arc), which form the ionized ring. The orbital plane of the ionized flows appears to be aligned with the plane of the circumnuclear disk (CND) derived by Güsten et al. (1987) from molecular line observations. If the fitted parameters from both the molecular lines (Güsten et al. 1987) and hydrogen recombination lines (Schwarz, Bregman, & van Gorkom 1989; Lacy, Achtermann, & Serabyn 1991; Roberts & Goss 1993; Vollmer & Duschl 2000; Liszt 2003; Zhao et al. 2009) are correct, the Keplerian motions dominate

the kinematics in the central parsec, and the plane of the main nuclear gaseous disk or the CND can be described by the mean orbital angles inferred above. Therefore, the orbital plane of the Eastern Arm ($i = 122^\circ \pm 5^\circ$ and $\Omega = -42^\circ \pm 11^\circ$) is nearly perpendicular to that of the CND. Figure 5 shows that the Eastern Arm flow runs into the CND from the northeast and the near side with respect to Sgr A* and collides with the Northern Arm flow in the region located south and behind Sgr A*.

3.2.2. Kinematics and Keplerian Motions

Figure 6 shows the spectra taken from 12 positions along each of the two ionized arms (shown in Figure 7). The observed spectra are compared with the radial velocities (green bars) derived from the Keplerian model. Note that we assume that the LSR velocity of the black hole, V_z , is 0 to within $\pm 5 \text{ km s}^{-1}$ (see discussion in Gillessen et al. (2009)). In general, the kinematics of the ionized gas follow the Keplerian orbital motion expected for the gravity of the SMBH ($4.2 \times 10^6 M_\odot$) at Sgr A*. Models of the enclosed mass (see Oh, Kim, & Figer 2009) show that the added enclosed mass starts to rise above the black hole’s mass starting at a radius of 0.3 pc ($7''$) from Sgr A*. At a radius of 1 pc ($24''$) from Sgr A*, near the tip of the minispiral seen in projection, the enclosed mass has nearly doubled, to about $8 \times 10^6 M_\odot$. Thus, for a given circular orbit with a radius of 1 pc the actual orbital velocity would be about 1.4 times higher than in a pure Keplerian approximation based on the mass of the black hole only. Apart from this effect, however, in several regions the radial velocities deviate even more significantly from those predicted from the Keplerian model.

3.2.3. Non-Keplerian Components

In the northwest end of the Eastern Arm near the location E17 (Figure 7a), a feature with radial velocities blue-shifted up to $\sim -200 \text{ km s}^{-1}$ shows a large deviation from the value of $\sim -100 \text{ km s}^{-1}$ predicted by the Keplerian model (Figure 6). From Paumard et al. (2006), we find that at least five massive stars—IRS 34W (Ofpe/WN9), IRS 34E (O9–9.5I), IRS 34NW (WN7), IRS 7SW (WN8), and IRS 3E (WC5/6)—are located in a region with a radius of $2.2''$ (indicated with a

large black circle near region E16 in Figure 7a), referred to as IRS 34 hereafter. Except for IRS 3E, which has no proper motion measurements, the proper motion data of the remaining four stars (Paumard et al. 2006) show that the stars all move southwest toward the ionized flow with a mean transverse velocity of 185 km s^{-1} ($\text{PA} = -134^\circ$) (as shown by an arrow in Figure 7a). We compare the radial velocity of the stars with the spectrum (inset c in Figure 7a) of the $\text{H}30\alpha$ line emission integrated from the IRS 34 region. The four stars besides IRS 3E show negative radial velocities in the range -150 km s^{-1} to -340 km s^{-1} , which is blue-shifted with respect to the peak velocity ($-146 \pm 4 \text{ km s}^{-1}$) of the broad ($\Delta V_{\text{FWHM}} = 160 \pm 8 \text{ km s}^{-1}$) spectral feature. The stellar winds from the Wolf-Rayet (WR) stars, evolved from early type O-stars with mass loss rate of $10^{-5} M_\odot \text{ yr}^{-1}$, appear to play a considerable role in the local kinematics. The strong stellar winds with velocities up to a few times 10^3 km s^{-1} could sweep up the ionized flow of the Eastern Arm, causing the peculiar velocities and the velocity gradient across the Eastern Arm at the northwest end. In fact, the motions of the filaments associated with the minispiral found by (Muzic et al. 2007) in infrared also suggest that a model of purely Keplerian motions of the gas is inadequate and the filaments could be due to the interaction of a fast wind with the minispiral.

In the tip region (Paumard, Maillard, & Morris 2004) with peculiar radial velocity that deviates significantly from the Keplerian velocity by $+150 \text{ km s}^{-1}$ (see panel E12 in Figure 6 and Figures 7a and 7b), there is also a WN8 star, IRS 9W. The radial velocity ($+140 \text{ km s}^{-1}$) of IRS 9W (Paumard, Maillard, & Morris 2004) is similar to that of ionized gas in the Eastern Arm. A comparison between the stellar transverse velocity of 215 km s^{-1} ($\text{PA} = 51^\circ$) and that of 304 km s^{-1} ($\text{PA} = -67^\circ$) predicted from the Keplerian model for the ionized flow at the location of IRS 9W suggests that the star does not follow the orbit of the Eastern Arm but may run across the Eastern Arm. Then, the high-radial velocity tip is likely to be the part of the gas in the ionized flow that is compressed and accelerated by the suspected strong wind from the IRS 9W star. The velocity discontinuity at the tip may be evidence for the presence of shocks in the region.

In addition, deviations from Keplerian motions are also observed in the Bar region. The non-Keplerian kinematics could be attributed both to the interaction between the ionized flows in the Northern and Eastern Arms and to the interaction of the ionized flows with the strong winds from the massive star clusters (the IRS 16 and IRS 13 clusters, for example). Well-organized polarized emission from magnetically aligned dust grains in the central parsec has been observed (Aitken et al. 1991, 1998; Glasse, Aitken & Roche 2003) suggesting a magnetic field strength of ≥ 2 mG. The corresponding energy densities in the magnetic fields are $\sim 1.6 \times 10^{-7}$ erg cm $^{-3}$, about twice larger than the thermal energy densities in the ionized gas ($T_e \sim 7 \times 10^3$ K and $n_e \sim 6 \times 10^4$ cm $^{-3}$ at IRS 1W) but two orders of magnitude less than the kinetic energy densities of the orbital motions. The magnetic fields frozen in the ionized streamers can not substantially alter the bulk motions of the gas but produce some cumulative magnetohydrodynamic (MHD) effects, such as orbital compression of the streamers as implied from the observed evidence for the convergence of field lines near Sgr A* (Glasse, Aitken & Roche 2003), and a helix of a natural morphology for a twisted magnetic field frozen into a plasma as observed in the Northern Arm (Yusef-Zadeh, Morris, & Ekers 1989; Zhao et al. 2009). In the region close to Sgr A*, the magnetic field probably becomes stronger, and the MHD effects may also play a significant role in the kinematics of the ionized flows.

An interesting anomalous feature located $\sim 2''$ northwest of Sgr A* can be seen in the integrated H30 α line intensity image (Figure 1). The radial-velocity image (Figure 2) indicates a large velocity gradient across this feature. The H30 α line spectrum (inset d in Figure 7) suggests that this feature is composed of at least three spectral components ($\geq 4\sigma$) at radial velocities of 160 km s $^{-1}$, 39 km s $^{-1}$, and -250 km s $^{-1}$, with a line width of ~ 50 km s $^{-1}$, a typical velocity width of the ionized gas in the minispiral arms. We note that there is a compact continuum source called ε (Yusef-Zadeh & Morris 1991), located $1''$ south of this anomalous feature, corresponding to a hole in the H30 α line images (e.g., Figure 7a and Figure 1). The source ε appears to be moving away from Sgr A* with a transverse velocity of 340 km s $^{-1}$ (Zhao et al. 2009), which has been interpreted as

the confluence of gravitationally focussed winds from the IRS 16 cluster (Wardle & Yusef-Zadeh 1992). The H30 α line property of the anomalous feature appears to be consistent with the focused-wind model, suggesting that the velocity components in the anomalous feature correspond to ionized blobs of the minispiral arms that have been blown out of the region ε due to the high pressure of the confluent stellar winds. However, given the closeness to Sgr A*, the detection of this feature needs to be further confirmed with higher-sensitivity observations.

Finally, in the locations with large angular offsets from Sgr A*, we noticed that the deviations in radial velocity from the Keplerian motions become significant for the Northern Arm (> 50 km s $^{-1}$), as shown in both Figures 6 and 7, while for the Eastern Arm, the deviations are less significant. The difference in the deviation of the observed radial velocities from the Keplerian velocities between the Northern and Eastern Arms suggests that the distribution of mass has a preference for the plane of the CND.

4. IONIZATION

4.1. The High-Density Gas in the Northern and Eastern Arms

We have shown that the H30 α lines observed with the SMA are excellent tracers of the high-density HII components in the Northern and Eastern Arms. Based on the observations and derived physical parameters, we can estimate the ionizing photon rate ($Q = n_e^2 \alpha_B V$) required for maintaining the ionizations in these regions for the case where the ionized gas is optically thick in the Lyman lines and where α_B is the ionization coefficient, and $V \approx \Omega D^2 L f_V$ is the volume of the ionized clumps. The ionizing photon rate can be written as

$$Q(H) = 1.31 \times 10^{45} \left[\frac{\alpha_B(T_e)}{\alpha_B(10^4 \text{ K})} \right] \left[\frac{\theta_B}{\text{arcsec}} \right]^2 \left[\frac{n_e}{10^4 \text{ cm}^{-3}} \right]^2 \left[\frac{D}{8 \text{ kpc}} \right]^2 \left[\frac{L f_V}{10^{-3} \text{ pc}} \right] \text{ phot. s}^{-1}, \quad (9)$$

where θ_B is the resolution of the observations (beam) with the assumption that source fills the beam; and $\alpha_B(T_e)$, the total recombination coefficient to excited levels, is given by Hummer & Seaton (1963). We calculated the values of the inferred ionizing photon flux required to maintain the ionization of the high-density components in the Northern and Eastern Arms as well as the HII mass (M_{HII}) and excitation parameter (U). On the assumption that the HII gas is internally ionized, the equivalent zero-age main sequence (ZAMS) star that is the ionizing source was estimated. The internal-ionization model requires at least ~ 25 O9-type stars embedded and roughly evenly distributed in both the Northern and Eastern Arms.

4.2. Ionizing Sources

We now examine where the known OB stars lie in the region. Figure 8 shows the 90 massive stars within a radius of $14''$ (0.5 pc) of Sgr A* identified by Paumard et al. (2006), among which 55 are OB stars. The 14 OB stars within a radius of $0.85''$ (0.03 pc) from Sgr A* are known as S stars. No obvious correlations appear between the location of 41 OB stars outside the radius of $0.85''$ and the distribution of the high-density HII gas in the minispiral arms on the projected plane. If we assume the average solid angle for the minispiral subtended to be Ω_s , then the required number of ionizing photons from the stars will be

$$Q^*(H) = Q(H)(4\pi/\Omega_s) . \quad (10)$$

If we assume $4\pi/\Omega_s \sim 10$, then 250 O9-type stars are required.

For the lower-density ($\bar{n}_e \sim 8 \times 10^2 \text{ cm}^{-3}$) component B, the ionized mass of $\sim 340 M_\odot$ is inferred, requiring a Lyman photon flux of $1.2 \times 10^{50} \text{ phot. s}^{-1}$ or ~ 100 O9-type stars (ZAMS). The 250 O9-type stars required for externally maintaining the ionization of the minispiral appear to be adequate in supplying the leaking Lyman photons from the minispiral to further ionize the lower density component that pervasively distributes in the region. Thus, the ionizing source to maintain the ionization state of Sgr A West is equivalent to ~ 250 O9-type stars (ZAMS). Considering that a substantial amount of the ionizing photons can

be produced by the hot evolved massive stars and the survey of Paumard et al. (2006) is incomplete, there may be enough stars to ionize the region.

The origin of the ionized minispiral arms and the birthplace of the massive stars found in the central parsec are still matters of conjecture. There are two possible models that may explain the presence of the massive stars near the SMBH. First, the stars formed elsewhere outside the central parsec and migrated to the Galactic center (Gerhard et al. 2001). However, compared with the age of the massive stars in the central parsec, it takes too long for stars to move to the center from several parsec away by losing their orbital energy through dynamical friction (Portegies Zwart, McMillan, & Gerhard 2003; Kim & Morris 2003). Alternatively, the stars may have formed near the SMBH by an unusual astrophysical process, such as the fragmentation of a massive disk rotating near the SMBH (Nayakshin, Guadra, & Springel 2007). Recent modeling and numerical simulations (Nayakshin, Guadra, & Springel 2007; Wardle & Yusef-Zadeh 2008; Bonnell & Rice 2008) show that for an infalling giant molecular cloud ($10^{4-5} M_\odot$) interacting with an SMBH ($1-3 \times 10^6 M_\odot$), a small fraction of the gas cloud becomes bounded to the black hole, forming an eccentric disk that quickly fragments to form stars. The tidal disruption of the self-gravitating clouds formed coherent spiral structures, i.e., the minispiral arms, which are ionized by the Lyman continuum photons produced from the newly formed massive stars.

5. SUMMARY AND CONCLUSION

We report results from the SMA observations of the H30 α line emission from the Galactic center at a resolution of $2''$. We measured the velocity profiles over the central $60''$ (2 parsec) region around Sgr A*. We carried out an analysis of the line and continuum data based on an isothermal homogeneous HII model for each component. With the constraints from the SMA H30 α data at 1.3 mm and the VLA H92 α data at 3.6 cm and continuum data at 1.3 cm, we determined the physical conditions of ionized gas in the Northern and Eastern Arms of the minispiral. We found that the typical electron density and kinetic temperature in the arms are $3-21 \times 10^4 \text{ cm}^{-3}$ and 5,000–13,000 K, re-

spectively. The highest density of $2.1 \times 10^5 \text{ cm}^{-3}$ occurs in the IRS 13 region. Higher temperatures up to 13,000 K are found in the Bar region. Both the H30 α and H92 α line profiles are broadened due to large velocity gradients along the line of sight and across the 2'' beam produced by the dynamical motions near Sgr A*. For the H92 α line, the line width due to pressure broadening appears to be comparable to that of thermal Doppler broadening, while for the H30 α lines, the pressure broadening is negligible.

Using the orbital parameters derived under the assumption that the three minispiral flows are in Keplerian motion, we calculated the 3D geometry of the minispiral structure. We showed that the ionized flows of Northern Arm and Western Arc are almost coplanar, and the plane of the ionized flow in the Eastern Arm is nearly perpendicular to that of the Northern Arm and Western Arc. The Eastern Arm flow collides with the Northern Arm flow in the Bar region located 0.1–0.2 parsec south of and behind Sgr A*. The high kinetic temperature gas in the Bar is probably due to the heat in shocks generated by collisions.

We compared observed radial velocities in the Northern and Eastern Arms with those computed from the Keplerian model with a mass of $4.2 \times 10^6 M_\odot$ of the SMBH at the position of Sgr A*. We show that Keplerian motion dominates the kinematics of the minispiral arms. In some regions, significant deviations exist between the observed radial velocities and those computed from the Keplerian model. Near the IRS 9W and IRS 34 regions, the observed peculiar velocities in the minispiral are likely a result of the ionized gas being swept up by the strong stellar winds from nearby WR stars.

A total Lyman continuum flux of $3 \times 10^{50} \text{ phot. s}^{-1}$ is needed to maintain the ionization of the gas in Sgr A West with an inferred ionized mass of $350 M_\odot$, which requires ~ 250 O9-type (ZAMS) stars. About 10% of the total number of UV photons is required to externally maintain the ionization of the high-density gas in the Northern and Eastern Arms, which have an ionized mass of $12 M_\odot$, less than 5% of the total ionized mass (Liszt 2003).

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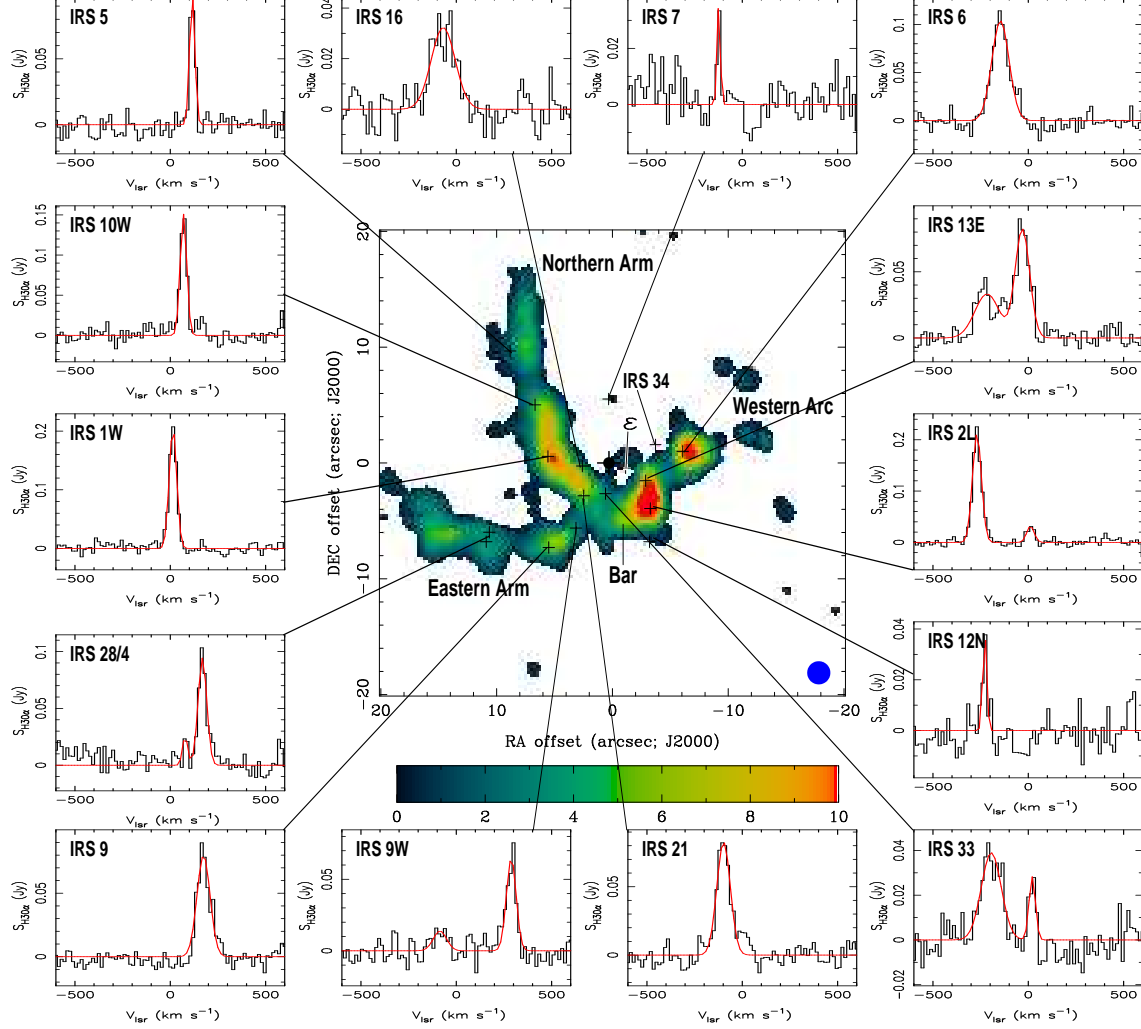


Fig. 1.— The image of the H30 α line intensity integrated from the velocity range between -360 and $+345$ km s^{-1} with the cutoff of 30 mJy beam^{-1} ($\sim 4\sigma$) in each of the channels. The color wedge shows the intensity in the range between 0 to $10 \text{ Jy beam}^{-1} \text{ km s}^{-1}$. The beam size ($\theta_{\text{FWHM}} = 2''$) is shown in the bottom right. The dot-cross at the coordinate origin marks the position of Sgr A*, where ($\alpha_{J2000} = 17^{\text{h}}45^{\text{m}}40.037^{\text{s}}$, $\delta_{J2000} = -29^{\circ}00'28.11''$) defines the origin of the coordinate system. The plus signs mark the positions from which the H30 α line spectra in the beam are taken and a few additional sources discussed in the text. The red curves are the Gaussian fits (multiple components in some cases) obtained from a least-squares analysis. The spectra are labeled with the names of their corresponding IRS sources.

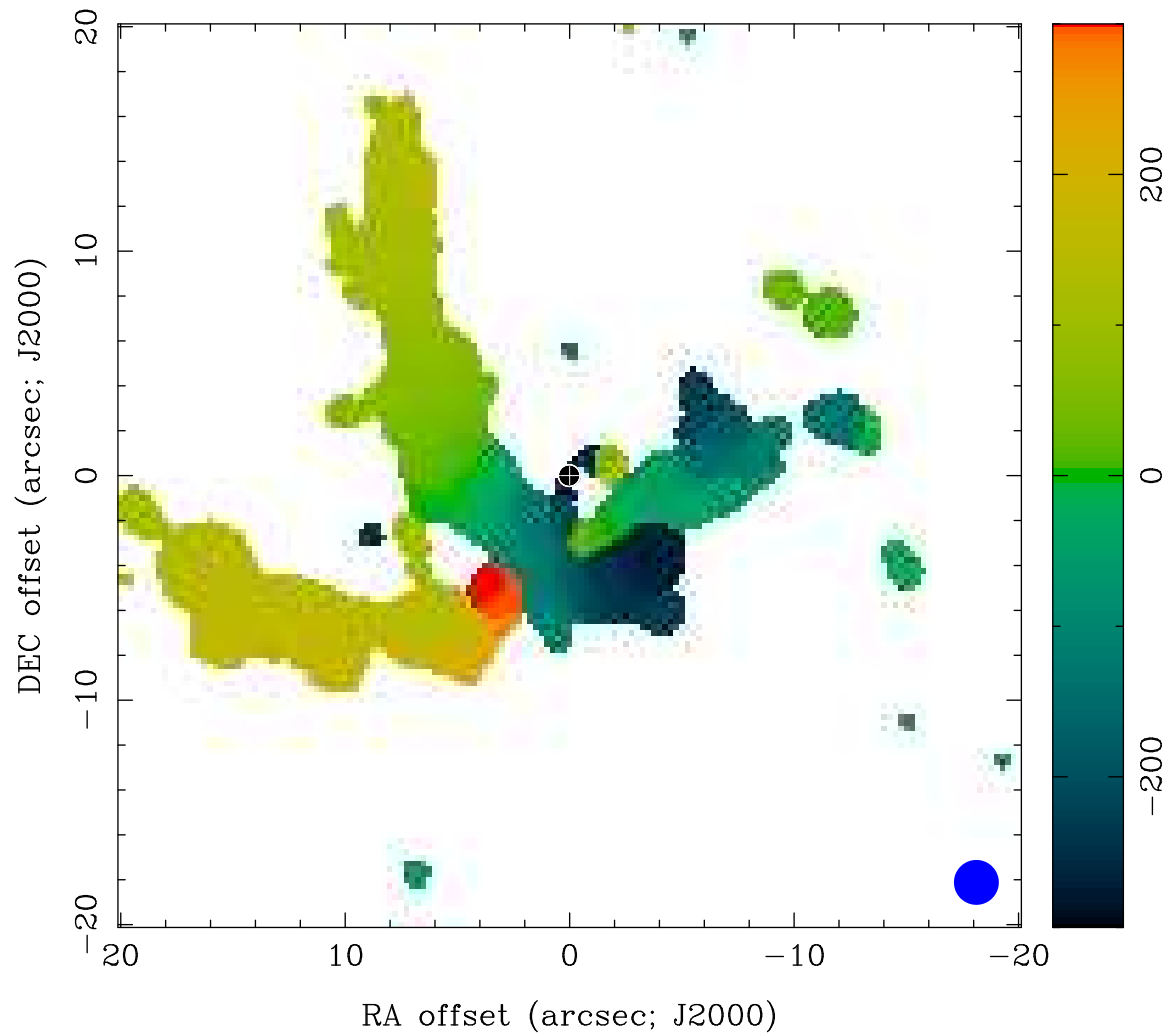


Fig. 2.— The distribution of the radial velocities determined from fitting the peak velocity of the H30 α line image cube observed with the SMA in the central 40'' region. The color scale denotes velocity from -360 to $+345$ km s $^{-1}$. The circular beam $\theta_{\text{FWHM}} = 2''$ is shown at the bottom right. The dot-cross at the coordinate origin marks the position of Sgr A*.

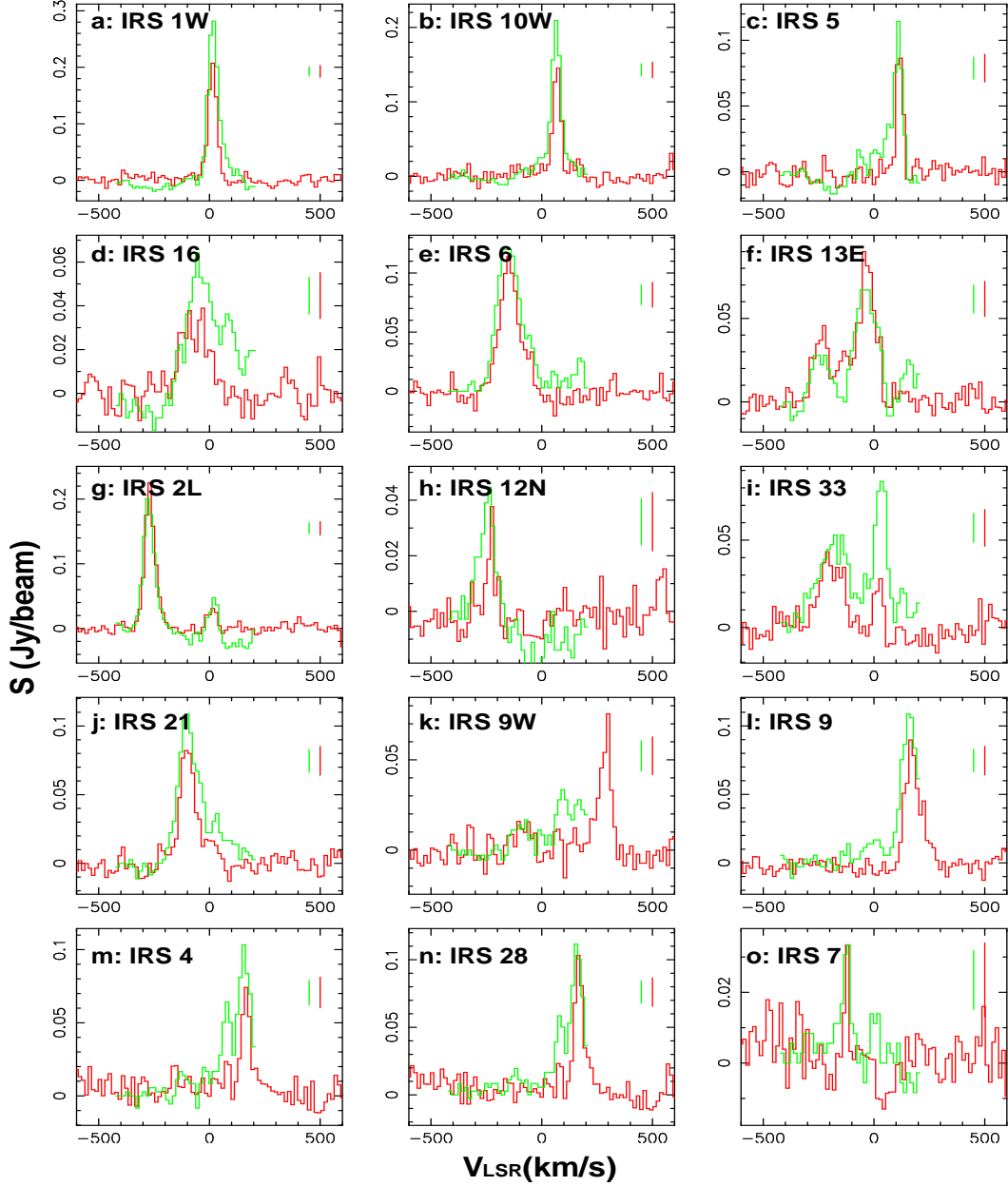


Fig. 3.— A comparison of the H30 α line spectra (red) with the VLA H92 α line spectra (green) taken from the selected IRS source regions with a resolution of $\theta_{\text{FWHM}} = 2''$. The VLA H92 α spectra have been multiplied by a factor of $\nu_{\text{H30}\alpha}/\nu_{\text{H92}\alpha} = 28$. The vertical bars indicate the uncertainty of 3σ .

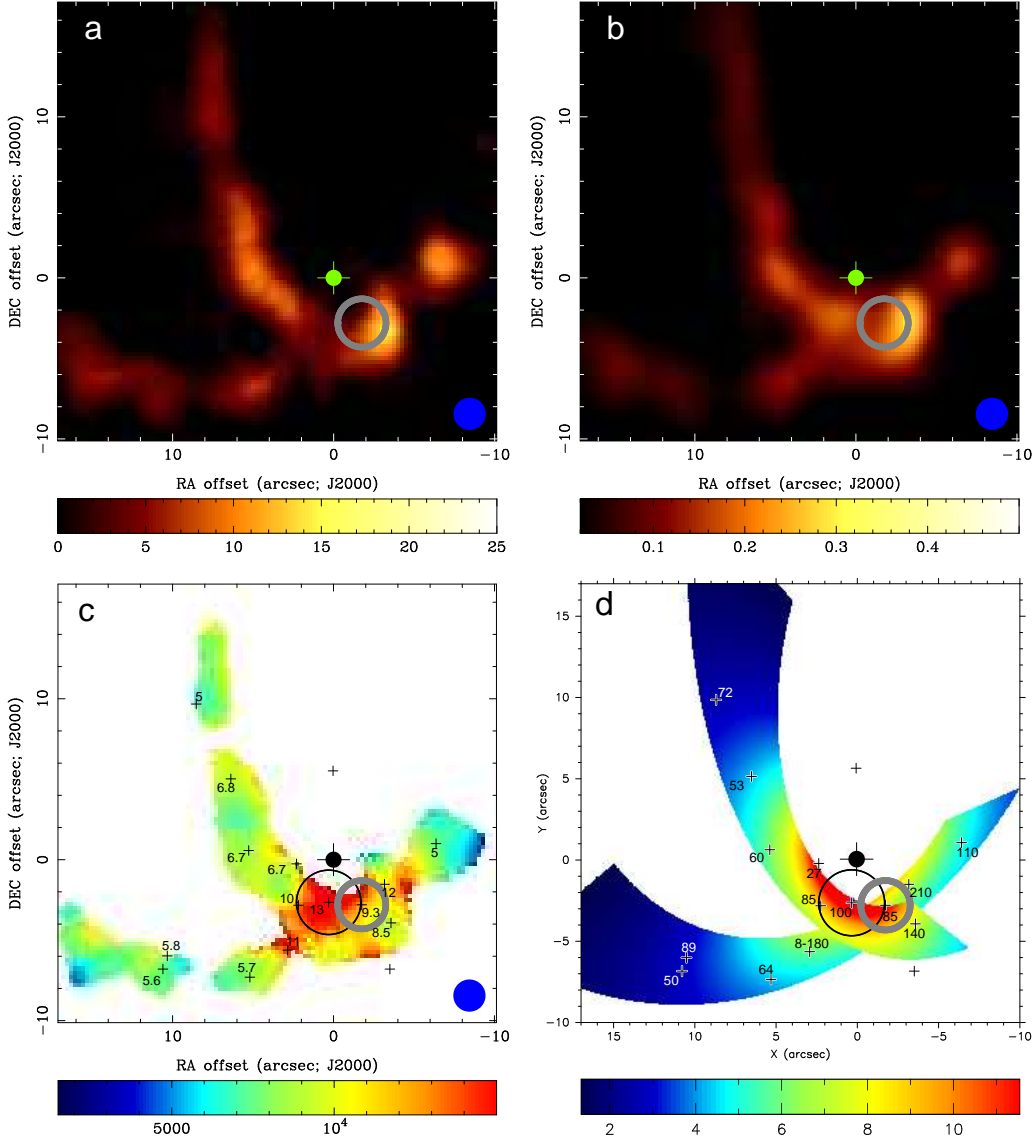


Fig. 4.— **a)** Integrated H30 α line emission observed at 1.3 mm with the SMA. **b)** Radio continuum image made from VLA observations at 1.3 cm with a resolution of $\theta_{\text{FWHM}} = 2''$. The radio emission associated with Sgr A* and the X-ray transient (J174540.0290031) has been removed, but the position of Sgr A* is marked by the dot-cross at the coordinate origin in each of the figures. **c)** Pseudo-color shows the LTE temperature (T_e^*) derived from images of H30 α line emission and radio continuum based on Equation 1. **d)** The image of the specific kinetic energy computed using the Keplerian orbital model described in Table 5 of Zhao et al. (2009). The color scale shows the specific kinetic energy in units of 10^{14} erg g $^{-1}$. In Figures 3c and 3d, the electron temperature (T_e in 10^3 K) and the electron density (n_e in 10^3 cm $^{-3}$) are shown at the positions of selected IR sources. These numbers are derived using the isothermal homogeneous HII model with the non-LTE effects. The results are tabulated in Table 3. The thick gray circles in (a),(b), (c) and (d) indicate the well-known minicavity region (Yusef-Zadeh, Morris, & Ekers 1989); the black circles in (c) and (d) indicate high-temperature regions.

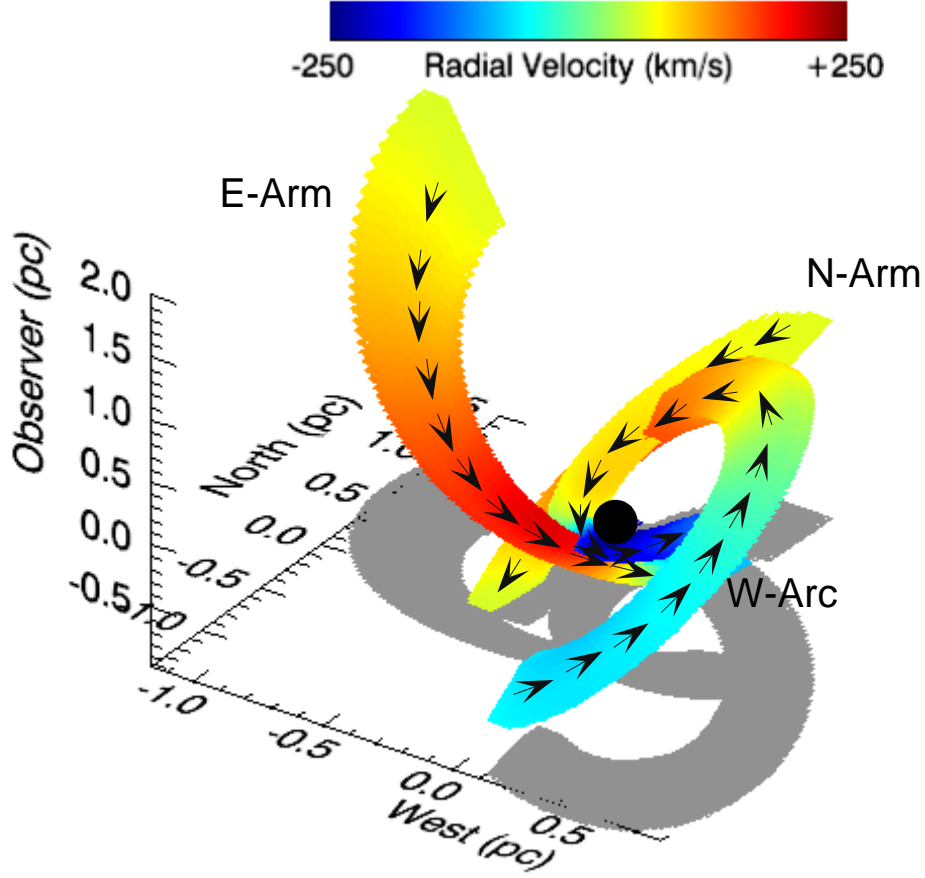


Fig. 5.— 3D view of the minispiral structure and velocity along each of the minispiral arms in the nucleus of the Galaxy centered at Sgr A* (black dot), which are computed using the Keplerian model described in Zhao et al. (2009). The orbital parameters were derived from fitting the VLA H92 α line data and are summarized in Table 5 of Zhao et al. (2009). The arrows indicate the direction of the ionized flows along the orbits. The color illustrates the radial velocity between -250 to $+250$ km s $^{-1}$ (color wedge at top). The gray shadow indicates the loci of the ionized flows projected on the sky plane.

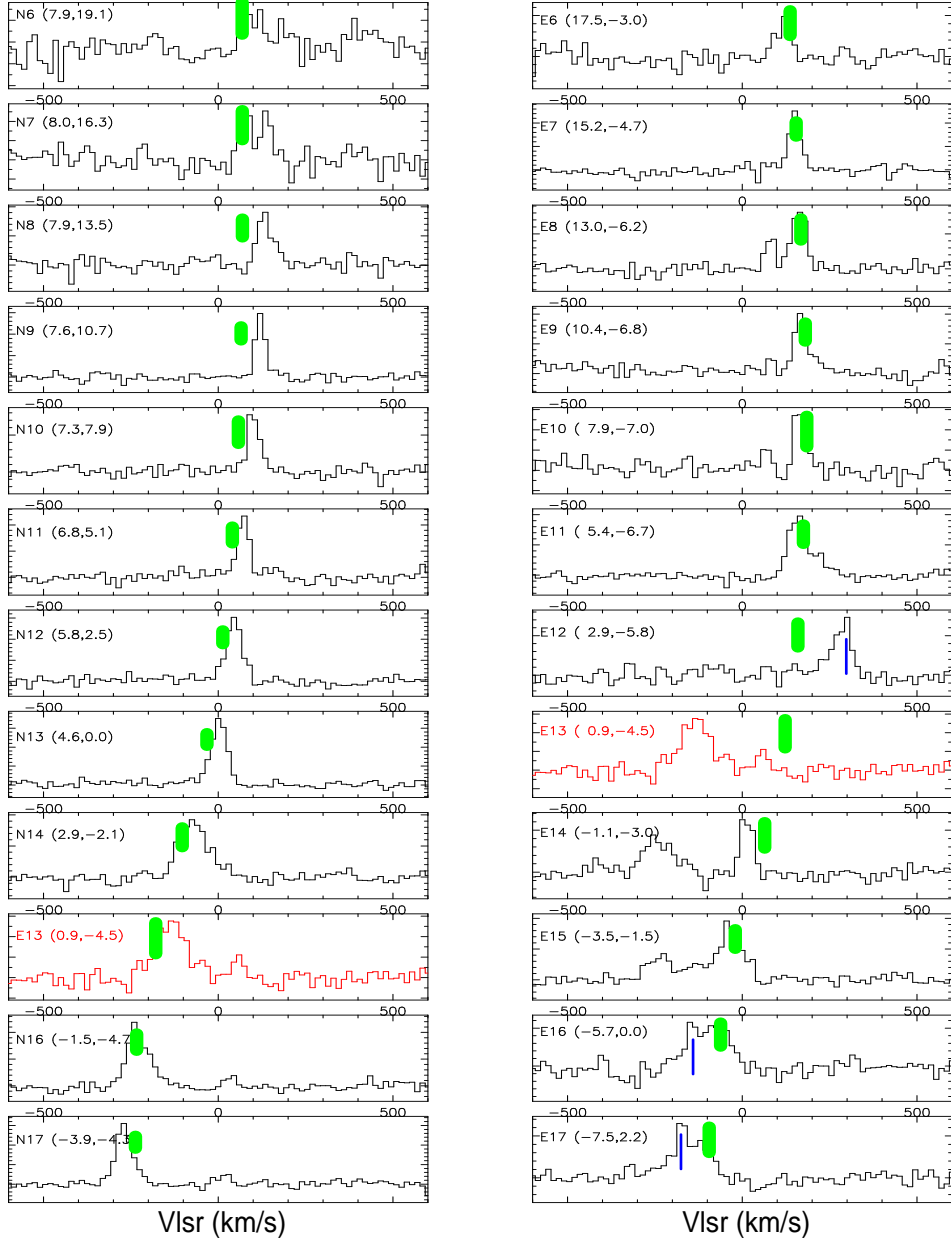


Fig. 6.— Left: Spectra of the H30 α line taken along the Northern Arm. Right: Spectra of the H30 α line taken along the Eastern Arm. The locations where the spectra were taken are shown in Figure 7. The green bars mark the radial velocities along the Northern and Eastern Arms computed using the Keplerian model with the orbital parameters given in Table 5 of Zhao et al. (2009), assuming that the mass of the SMBH is $4.2 \times 10^6 M_{\odot}$ and its radial velocity $V_z = 0 \text{ km s}^{-1}$ with respect to the LSR. The thin blue bars (in panels E12, E16, and E17) indicate peculiar velocities possible due to interactions with the strong stellar winds from the nearby WR stars (see Section 3.2.3). Note that the spectrum marked “E13” and shown in red in both columns comes from a common position, with the appropriate Keplerian velocity marked for each arm.

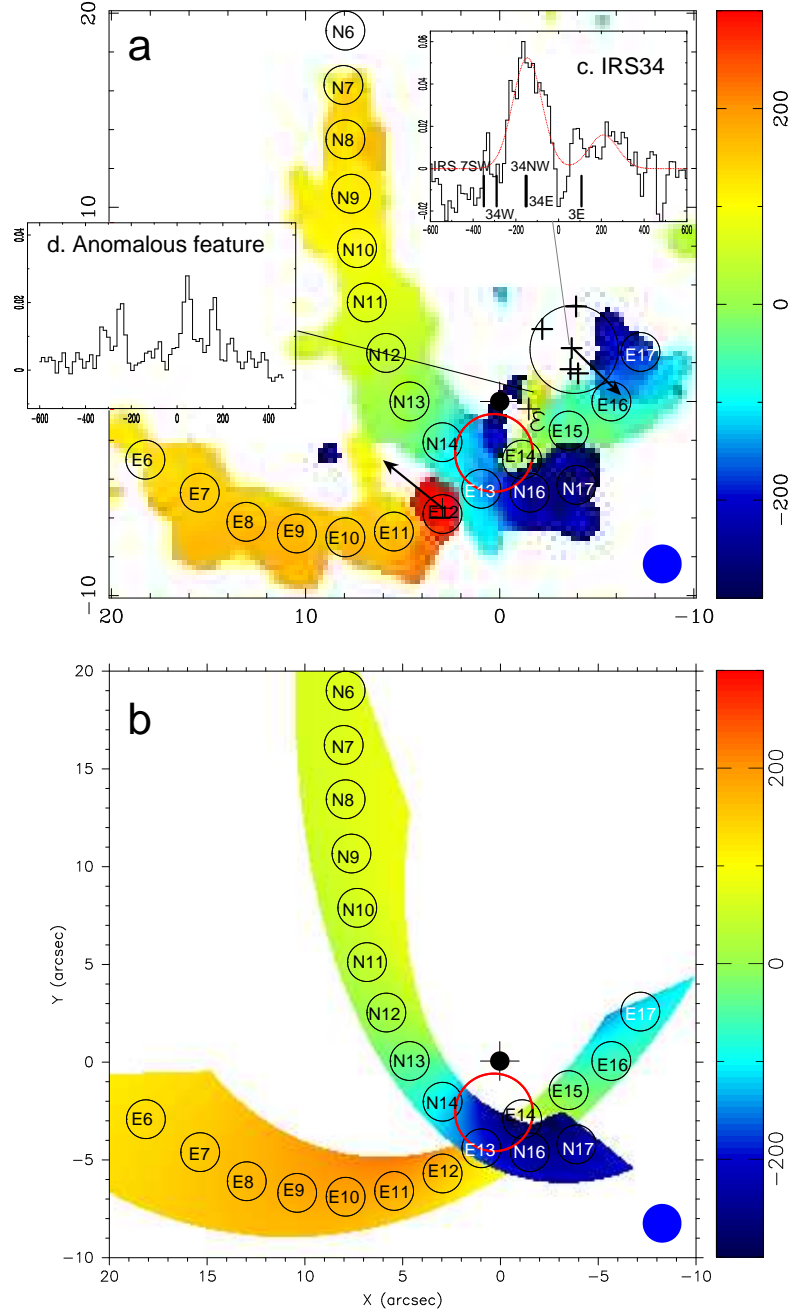


Fig. 7.— A comparison between the observed radial velocities (a) and those computed from the Keplerian model (b). The black circles along the two spiral arms mark the locations of the H30 α spectra displayed in Figure 6. The red circle indicates the high temperature region. The dot-cross marks the position of Sgr A* at the coordinate origin. The inset (c) is the line spectrum of H30 α toward the IRS 34 region indicated with a large black circle, which includes the five massive stars IRS 7SW (WN8), IRS 34W (Ofpe/WN9), IRS 34E (O9–9.5I), IRS 34NW (WN7), and IRS 3E (WC5/6) (Paumard et al. 2006). The mean proper motion of the first four stars (and a fifth star without proper motion) is indicated with an arrow near the E16 region in (a). The radial velocities of the five stars are marked with bars in inset (c). The proper motion of the IRS 9W star (WN8) from Paumard et al. (2006) is indicated with an arrow near the E12 region in (a). The inset (d) shows the H30 α line spectrum toward an anomalous feature ($-1.5''$, $0.5''$) northwest of Sgr A* and $1''$ north of ϵ .

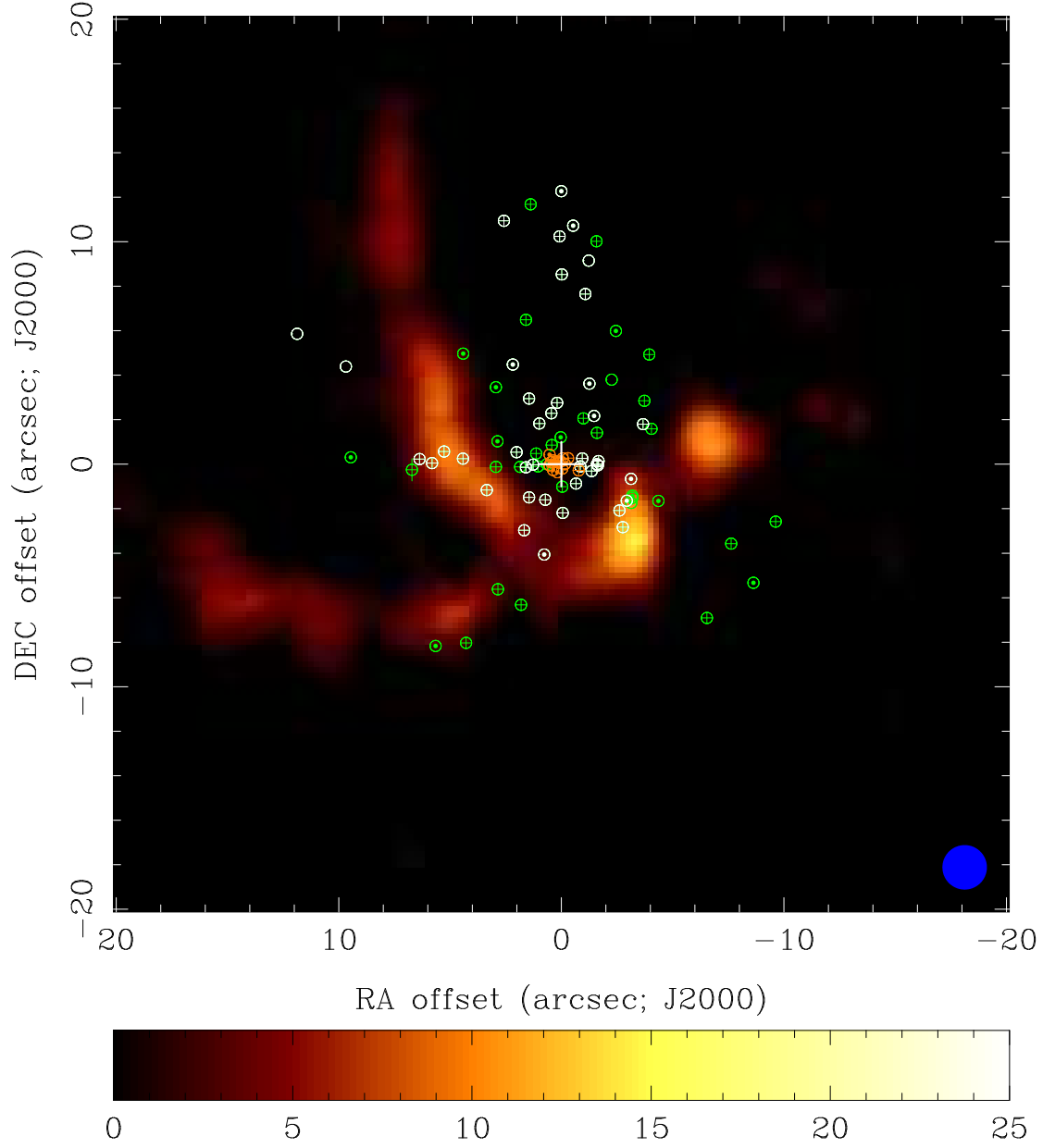


Fig. 8.— The 90 massive stars of Paumard et al. (2006) are overlaid on the H30 α line image. Among the 90 massive stars, 14 OB stars (orange), known as S stars, are located within a radius of 0.85'' from Sgr A* (white plus sign), 41 OB stars (white) are outside the S cluster, and the rest are Ofpe/WN9 WR stars (green). The cross- and dot-circles represent the stars with proper motion data, and the open circles indicate stars with no proper motion data.

TABLE 1
SUMMARY OF THE H30 α LINE OBSERVATIONS AT $\nu = 231.901$ GHz

Date	t_{obs} on Sgr A* (hr)	T_{sys} (K)	Array ^a	Calibrators
2006-04-08	1.79	96	C	3C273 ^b , Callisto ^c
2006-05-28	3.99	173	E	3C279 ^b , Uranus ^b , J1924-292 ^b , Callisto ^{b,c}
2006-07-17	2.19	146	V	3C279 ^b , J1924-292 ^b , Callisto ^c
2007-03-31	3.67	108	C	3C273 ^b , Callisto ^c
2007-04-01	3.35	126	C	3C273 ^b , Callisto ^c
2007-04-03	1.99	93	C	3C273 ^b , Callisto ^c
2007-04-04	1.93	103	C	3C273 ^b , Callisto ^c
2007-08-14	3.63	115	E	3C279 ^b , NRAO 530 ^c
2007-08-20	4.94	95	E	3C454.3 ^b , NRAO 530 ^c
2007-08-31	4.19	134	E	3C454.3 ^b , NRAO 530 ^c
2008-05-05	2.67	156	C	3C273 ^b , Ganymede ^c

^aArray configuration: C = Compact; E = Extended; V = Very Extended

^bCalibrator: bandpass

^cCalibrator: flux density scale

TABLE 2
THE H30 α & H92 α LINE MEASUREMENTS

Region	$\Delta\alpha$ (arcsec)	$\Delta\delta$ (arcsec)	Radio Properties in IRS Regions					$V_{\text{LSR}}^{\text{a}}$ (km s $^{-1}$)	$\Delta V_{\text{FWHM}}^{\text{a}}$ (km s $^{-1}$)
			$S_{\text{H30}\alpha}$ (Jy beam $^{-1}$)	$S_{22\text{GHz}}$ (Jy beam $^{-1}$)	$\frac{S_{\text{H30}\alpha}}{S_{22\text{GHz}}}$	$\frac{S_{\text{H92}\alpha}}{S_{22\text{GHz}}}$	$\frac{S_{\text{H30}\alpha}}{S_{\text{H92}\alpha}}$		
IRS 1W	5.27	0.57	0.21 \pm 0.008	0.16 \pm 0.008	1.3 \pm 0.08	0.06 \pm 0.005	21 \pm 1	14 \pm 1	46 \pm 2
IRS 10W	6.38	5.02	0.15 \pm 0.018	0.10 \pm 0.003	1.5 \pm 0.19	0.07 \pm 0.004	21 \pm 1	69 \pm 2	39 \pm 5
IRS 5	8.53	9.63	0.10 \pm 0.019	0.04 \pm 0.006	2.5 \pm 0.6	0.1 \pm 0.02	25 \pm 2	115 \pm 3	33 \pm 8
IRS 16	2.30	-0.26	0.03 \pm 0.005	0.08 \pm 0.015	0.4 \pm 0.1	0.026 \pm 0.008	15 \pm 4	-69 \pm 10	141 \pm 27
IRS 6	-6.37	1.00	0.12 \pm 0.011	0.15 \pm 0.006	0.80 \pm 0.08	0.034 \pm 0.003	23 \pm 2	-146 \pm 5	109 \pm 13
IRS13E	-3.17	-1.53	0.08 \pm 0.006	0.25 \pm 0.030	0.32 \pm 0.05	0.010 \pm 0.002	32 \pm 4	-31 \pm 3	86 \pm 8
IRS 2L	-3.57	-3.93	0.21 \pm 0.011	0.23 \pm 0.030	0.91 \pm 0.13	0.031 \pm 0.005	29 \pm 1	-269 \pm 1	55 \pm 3
IRS 12N	-3.50	-6.80	0.04 \pm 0.010	0.04 \pm 0.002	0.76 \pm 0.22	0.036 \pm 0.013	21 \pm 6	-229 \pm 7	40 \pm 17
IRS 33	0.30	-2.65	0.04 \pm 0.008	0.19 \pm 0.024	0.21 \pm 0.05	0.009 \pm 0.003	24 \pm 5	-195 \pm 12	126 \pm 29
IRS 21	2.18	-2.83	0.08 \pm 0.009	0.16 \pm 0.018	0.50 \pm 0.08	0.023 \pm 0.005	22 \pm 3	-99 \pm 4	80 \pm 11
IRS 9W	2.85	-5.62	0.06 \pm 0.007	0.11 \pm 0.011	0.57 \pm 0.09	... ^b	... ^b	288 \pm 4	60 \pm 8
IRS 9	5.20	-7.30	0.08 \pm 0.009	0.09 \pm 0.009	0.85 \pm 0.13	0.043 \pm 0.008	20 \pm 2	174 \pm 5	84 \pm 11
IRS 4	10.33	-5.98	0.07 \pm 0.012	0.05 \pm 0.006	1.4 \pm 0.3	0.07 \pm 0.02	20 \pm 3	164 \pm 4	47 \pm 10
IRS 28	10.60	-6.80	0.10 \pm 0.010	0.07 \pm 0.006	1.5 \pm 0.2	0.06 \pm 0.01	25 \pm 2	168 \pm 4	49 \pm 22
IRS 7	0.03	5.52	0.03 \pm 0.010	0.15 \pm 0.001	2.0 \pm 0.6	0.07 \pm 0.03	28 \pm 8	-124 \pm 10	29 \pm 22
Minicavity ^g	-1.75	-2.80	0.08 \pm 0.006	0.18 \pm 0.020	0.45 \pm 0.07	0.016 \pm 0.002	29 \pm 2	6 \pm 2	41 \pm 4

Component	Radio Properties of Sgr A West				
	$\int S_{\text{H30}\alpha} dV$ (Jy km s $^{-1}$)	$\int S_{\text{H92}\alpha} dV$ (Jy km s $^{-1}$)	$S_{22\text{GHz}}$ (Jy)	Area (arcmin 2)	Area (pc 2)
Sgr A West total ^c	1700 \pm 300 ^d	60 \pm 7	16.3 \pm 0.2	1.0 \pm 0.1	5.5 \pm 0.5
A ^e	300 \pm 20	19 \pm 2	7.1 \pm 0.1	0.10 \pm 0.01	0.55 \pm 0.05
B ^f	1400 \pm 300	41 \pm 7	9.2 \pm 0.2	1.0 \pm 0.1	5.5 \pm 0.5

^aDetermined from the H30 α line.

^bNo H92 α data at the line velocity.

^cThe overall Sgr A West region (80'' \times 45'').

^dThe line flux determined from the IRAM 30m observation.

^eThe bright H30 α line emission region (the Northern and Eastern Arms in Fig. 2).

^fThe residual of Sgr A West after taking out the contribution from component A.

^gThe line and continuum intensities listed in this row reflect averaged values of the emission from both the rim and the void of the minicavity in the 2'' beam that is similar to the diameter of the minicavity.

TABLE 3
PHYSICAL CONDITIONS OF IONIZED GAS IN IRS REGIONS

Region	T_e (10^3 K)	n_e (10^4 cm $^{-3}$)	EM (10^7 cm $^{-6}$ pc)	$\tau_C(22\text{GHz})$ (10^{-3})	$\left(\frac{T_e}{T_e^*}\right)^\dagger$	$\left(\frac{\Delta V_t}{\Delta V_{\text{FWHM}}}\right)$	$\left(\frac{\Delta V_P^{\text{H92}\alpha}}{\Delta V_D^{\text{TH}}}\right)^\ddagger$
IRS 1W	6.7 \pm 0.5	6.0 \pm 0.4	1.8 \pm 0.1	14. \pm 1.1	0.71	0.88	0.82
IRS 10W	6.8 \pm 0.6	5.3 \pm 0.4	1.1 \pm 0.1	8.5 \pm 1.0	0.71	0.85	0.72
IRS 5	5.0 \pm 1.1	7.2 \pm 1.2	0.4 \pm 0.1	4.7 \pm 1.3	0.68	0.78	1.2
IRS 16	6.7 \pm 2.0	2.7 \pm 0.7	0.8 \pm 0.2	6.4 \pm 2.7	0.68	0.99	0.37
IRS 6	5.0 \pm 0.6	11. \pm 1.1	1.6 \pm 0.1	18. \pm 2.3	0.71	0.96	1.8
IRS 13E	12. \pm 2.5	21. \pm 3.6	3.2 \pm 0.5	12. \pm 1.9	0.84	0.84	2.0
IRS 2L	8.5 \pm 1.3	14. \pm 1.4	2.7 \pm 0.4	16. \pm 2.9	0.78	0.80	1.7
IRS 12N	12. \pm 4.7	2.8 \pm 1.0	0.5 \pm 0.2	1.8 \pm 0.6	0.77	0.80	0.27
IRS 33	13. \pm 3.3	10. \pm 3.0	2.6 \pm 0.3	8.5 \pm 1.8	0.81	0.97	0.92
IRS 21	10. \pm 1.8	8.5 \pm 1.3	2.0 \pm 0.3	9.1 \pm 2.0	0.78	0.94	0.92
IRS 9W	11.–12.	0.8–18	1.3–1.3	5.5–4.9	0.74–0.83	0.76–0.93	1.7–0.08
IRS 9	5.7 \pm 0.9	6.4 \pm 0.8	1.0 \pm 0.2	9.7 \pm 1.8	0.69	0.97	0.97
IRS 4	5.8 \pm 1.3	5.0 \pm 1.0	0.5 \pm 0.1	4.8 \pm 1.4	0.68	0.91	0.75
IRS 28	5.6 \pm 0.8	8.9 \pm 1.0	0.7 \pm 0.1	6.8 \pm 1.1	0.71	0.86	1.4
IRS 7	7.0 \pm 2.5	8.8 \pm 2.7	0.2 \pm 0.1	1.2 \pm 0.5	0.74	0.64	1.2
Minicavity	9.3 \pm 0.9	8.5 \pm 0.7	2.0 \pm 0.3	10. \pm 2.0	0.77	0.78	0.96

$^\dagger \frac{T_e}{T_e^*} \approx \left[b_{30} \left(1 - \frac{1}{2} \beta_{30} \tau_C(232 \text{ GHz}) \right) \right]^{0.87} \approx b_{30}^{0.87}$ for a very small value of $\tau_C(232 \text{ GHz})$ in a region at the Galactic center.

$^\ddagger \Delta V_D^{\text{TH}} = \sqrt{\frac{8 \ln(2) k T_e}{M_H}}$ is the thermal Doppler line width.